

Scalar Field as Dark Matter and Machian Cosmological Solution in the Generalized Scalar-Tensor Theory of Gravitation

A. Miyazaki*

Department of Physics, Loyola University, New Orleans, LA 70118

and

Faculty of Economics, Nagasaki Prefectural University

Sasebo, Nagasaki 858-8580, Japan

Abstract

The Machian cosmological solution satisfying $\phi = O(\rho/\omega)$ in the generalized scalar-tensor theory of gravitation with the varying cosmological constant is summarized. The scalar field ϕ with the exponential potential is introduced as dark matter and the barotropic evolution of matter in the universe is discussed. As the universe expands, the coefficient γ of the equation of state approaches to $-1/3$ and the coupling function $\omega(\phi)$ diverges to $-\infty$.

PACS numbers: 04.50.+h, 98.80.-k

It is believed that the field equations of the Brans-Dicke theory [1] coincide with those of the Einstein theory with the same energy-momentum tensor when the coupling parameter ω goes to the infinity,

*Email: miyazaki@loyno.edu, miyazaki@nagasaki-pu.ac.jp

but this is generally not true [2]–[4]. We proposed the postulate in the Machian point of view: *The scalar field of a proper cosmological solution should have the asymptotic form $\phi = O(\rho/\omega)$ when ω is large enough and should converge to zero in the continuous limit $\rho/\omega \rightarrow 0$ (the Machian solution)* [4]. We proved that the Machian solution necessarily satisfies the relation $GM/R = \text{const}$ for the homogeneous and isotropic universe [5].

We extended the Machian solution to the case of a perfect fluid with *negative* pressure [6] (see also [15]). This solution shows that the time-variation of the gravitational constant vanishes in the limit of $\gamma = -1/3$ and suggests that the coupling parameter ω of the Brans–Dicke scalar field should vary in time. We investigated various possibilities of the coupling function $\omega(\phi)$ in the generalized scalar–tensor theory of gravitation [7]–[9] and obtained $\omega(\phi) = \eta/(\xi - 2)$ to support the reasonable Machian solution [10]. In this cosmological model, the coefficient of the expansion parameter and the scalar field ϕ converge to constants respectively as the parameter $\xi \rightarrow 2(\gamma \rightarrow -1/3)$. When $\xi \rightarrow 2$, the coupling function $\omega(\phi)$ diverges to the minus infinity and the gravitational constant approaches dynamically to the constant G_∞ . Moreover, if we require the particular form of the cosmological constant, the similar Machian cosmological solution exists in the generalized scalar–tensor theory of gravitation and the universe shows the slowly accelerating expansion at present [11], which is compatible with the recent measurements [12] of the distances to type Ia supernovae.

In this short note, we discuss the scalar field φ as dark matter and the evolution of matter in the universe after summarizing the Machian cosmological solution in the (modified) generalized scalar–tensor theory of gravitation with the varying cosmological constant. Let us start

the following action

$$S = \int d^4x \sqrt{-g} \left\{ -\phi [R + 2 \lambda(\phi, \phi_{,\mu} \phi^{,\mu})] + 16\pi L_m - \frac{\omega(\phi)}{\phi} g^{\mu\nu} \phi_{,\mu} \phi_{,\nu} \right\}, \quad (1)$$

where $\omega(\phi)$ is an arbitrary coupling function and the cosmological constant is given as

$$\lambda(\phi, \phi_{,\mu} \phi^{,\mu}) \equiv \frac{1}{2} \lambda_1 \phi_{,\mu} \phi^{,\mu} / \phi^2. \quad (2)$$

The action Eq. (1) is equivalent to that of the generalized scalar-tensor theory without the cosmological constant after replacing $\omega(\phi) \rightarrow [\omega(\phi) + \lambda_1]$.

The energy-momentum tensor $T_{\mu\nu}$ of matter for the perfect fluid is in general described as

$$T_{\mu\nu} = -p g_{\mu\nu} - (\rho + p) u_\mu u_\nu, \quad (3)$$

and the conservation law of the energy-momentum $T^{\mu\nu}_{;\nu} = 0$ gives the equation of continuity $\dot{\rho} + 3(\dot{a}/a)(\rho + p) = 0$. Assuming the barotropic equation of state

$$p(t) = \gamma \rho(t), \quad -1 \leq \gamma \leq 1/3, \quad (4)$$

we get

$$\rho(t) a^n(t) = \text{const}, \quad (5)$$

where $n = 3(\gamma + 1)$.

We are interested in the closed model ($k = +1$) and the line element for the Friedmann-Robertson-Walker metric is

$$ds^2 = -dt^2 + a^2(t) [d\chi^2 + \sin^2\chi (d\theta^2 + \sin^2\theta d\varphi^2)]. \quad (6)$$

We obtain the independent field equations

$$\begin{aligned} \frac{3}{a^2} (\dot{a}^2 + 1) &= \frac{[\omega(\phi) + \lambda_1]}{2} \left(\frac{\dot{\phi}}{\phi} \right)^2 + \frac{\ddot{\phi}}{\phi} + \\ &+ \frac{8\pi\rho}{\phi} - \frac{8\pi(\rho - 3p)}{3 + 2[\omega(\phi) + \lambda_1]} \frac{1}{\phi} \end{aligned} \quad (7)$$

and

$$\ddot{\phi} + 3 \frac{\dot{a}}{a} \dot{\phi} = \frac{1}{3 + 2 [\omega(\phi) + \lambda_1]} \left[8 \pi (\rho - 3 p) - \frac{d\omega(\phi)}{d\phi} \dot{\phi}^2 \right]. \quad (8)$$

We require the particular coupling function

$$\omega(\phi) \equiv \frac{\eta}{\xi - 2}, \quad (9)$$

where $\xi = 1 - 3 \gamma$ or $\xi = 4 - n$, η is a constant, and this gives $d\omega/d\phi =$

0. We introduce another scalar function $\Phi(t)$ by

$$\phi(t) = \frac{8 \pi}{3 + 2 [\omega(\phi) + \lambda_1]} \Phi(t) \quad (10)$$

for the Machian solution satisfying $\phi = O(\rho/\omega)$ and the same relation between $\dot{\phi}(t)$ and $\dot{\Phi}(t)$, which is guaranteed for the very slow time-variation or the large value of ω . From Eqs.(7) and (8), we find that the expansion parameter has from

$$a(t) \equiv A(\omega) \alpha(t), \quad (11)$$

where

$$\frac{3}{A^2(\omega)} = \left| \frac{\omega(\phi)}{2} + B \right|. \quad (12)$$

Thus the field equations has a solution

$$\Phi(t) = \xi \rho(t) t^2, \quad \alpha(t) = bt, \quad (13)$$

with

$$\xi = 1/(\xi - 2), \quad (14)$$

$$b = \begin{cases} (4 - \xi^2)^{-1/2}, & \text{for } \omega/2 + B < 0 \text{ and } 0 \leq \xi < 2 \\ (\xi^2 - 4)^{-1/2}, & \text{for } \omega/2 + B > 0 \text{ and } 2 < \xi \leq 4, \end{cases} \quad (15)$$

and

$$B = \frac{\lambda_1}{2} - \frac{3}{(\xi - 2)(\xi + 2)}. \quad (16)$$

There is a discontinuity at $\xi = 2$ and we restrict to the range $0 \leq \xi < 2$ owing to the evolutionary continuity from $\xi = 0$ (the radiation era)

and $\xi = 1$ (the dust-dominated era).

Both the coupling function $\omega(\phi) = \eta/(\xi - 2)$ for $0 \leq \xi < 2$ and the constraint $\omega(\phi)/2 + B < 0$ lead necessarily to $\eta > 0$ and $\omega < 0$. We require $\eta > 3$ to avoid the singularity ($\omega < -3/2$). We find the expansion parameter $a(t)$ and its asymptotic form when $\xi \rightarrow 2$

$$a(t) = [6/f(\xi)]^{1/2} t \cong \sqrt{3/(2\eta - 3)} t, \quad (17)$$

where the quadratic function $f(\xi) \equiv \lambda_1(\xi - 2)(\xi + 2) + \eta(\xi + 2) - 6$. The value $\eta = 3$ gives the expansion parameter $a(t) = t$ (light velocity) when $\xi \rightarrow 2$. If we require $\lambda_1 < 0$ and $0 \leq -\eta/\lambda_1 < 2$, the universe (the closed space) shows the slowly accelerating expansion for the period $-\eta/\lambda_1 \leq \xi < 2$.

We find the scalar field $\phi(t)$ and its asymptotic form when $\xi \rightarrow 2$

$$\phi(t) = \frac{8\pi\rho(t)t^2}{(3 + \lambda_1)(\xi - 2) + 2\eta} \cong \frac{4\pi\rho(t)t^2}{\eta} \rightarrow \text{const}, \quad (18)$$

which converges to a definite and finite constant in the limit. As the parameter $\xi \rightarrow 2$, the coupling function $\omega(\phi)$ diverges to $-\infty$ and the gravitational constant approaches dynamically to G_∞ . The cosmological constant $\lambda(t)$ decreases rapidly in proportion to t^{-2} as the universe expands and converges to zero when $\xi \rightarrow 2$:

$$\lambda(t) = \frac{\lambda_1}{2} \left(\frac{\dot{\phi}}{\phi} \right)^2 \propto \frac{(\xi - 2)^2}{t^2} \rightarrow 0. \quad (19)$$

When $\xi \rightarrow 2$ ($t \rightarrow +\infty$), the abnormal term $\omega(\phi)a^2(\dot{\phi}/\phi)^2$ vanishes and the generalized scalar-tensor theory of gravitation reproduces the correspondent solution of general relativity with the same energy-momentum tensor, that is, the Friedmann universe with $\ddot{a}(t) = 0$, $\lambda(t) = 0$, and $p = -\rho/3$.

Let us introduce a scalar field φ , as dark matter to complete the physical evolution of matter in the universe, the Lagrangian of which is

given as

$$L_\varphi = - (1/2) g^{\mu\nu} \varphi_{,\mu} \varphi_{,\nu} + V(\varphi). \quad (20)$$

The energy-momentum tensor $T_{\mu\nu}^\varphi$ derived from this Lagrangian is

$$T_{\mu\nu}^\varphi = -\varphi_{,\mu} \varphi_{,\nu} + (1/2) g_{\mu\nu} \varphi_{,\lambda} \varphi^{,\lambda} + g_{\mu\nu} V(\varphi), \quad (21)$$

and the field equation of the scalar field $\varphi(t)$ is

$$\ddot{\varphi} + 3 \frac{\dot{a}}{a} \dot{\varphi} + \frac{dV}{d\varphi} = 0. \quad (22)$$

Let us require a scalar potential $V(\varphi)$ (see [13]):

$$V(\varphi(a)) \equiv V_0/a^2 > 0. \quad (23)$$

Taking this potential and the expansion parameter $a(t) = Abt$ into account, we find a solution of Eq. (22)

$$\varphi(a) = (\sqrt{V_0}/Ab) \ln a, \quad (24)$$

and thus we obtain the exponential potential of the scalar field φ

$$V(\varphi) = V_0 \exp[-(Ab/\sqrt{V_0})\varphi]. \quad (25)$$

The density of the scalar field φ is given as

$$\rho_\varphi = (1/2) \dot{\varphi}^2 + V(\varphi) = (3/2) (V_0/a^2) \propto t^{-2}, \quad (26)$$

and the pressure is

$$p_\varphi = (1/2) \dot{\varphi}^2 - V(\varphi) = -(1/2) (V_0/a^2). \quad (27)$$

Thus, we get the relation $p_\varphi = -(1/3) \rho_\varphi$ and may regard the scalar field φ as a perfect fluid with the barotropic equation of state Eq. (4).

It should be noted that the above Machian cosmological solution produces no changes by introducing the scalar field φ . Let us suppose that the total matter of the universe consists of the radiation (density ρ_r , $\gamma = 1/3$), the dust-matter (density ρ_m , $\gamma = 0$), and the scalar field φ (density ρ_φ , $\gamma = -1/3$) for simplicity:

$$\rho = \rho_r + \rho_m + \rho_\varphi. \quad (28)$$

The densities ρ_r , ρ_m , and ρ_φ satisfy the conservation laws $\rho_r a^4 = \rho_{r0} a_0^4$, $\rho_m a^3 = \rho_{m0} a_0^3$, and $\rho_\varphi a^2 = \rho_{\varphi 0} a_0^2$ respectively, where a subscript 0 denotes

the present value. The total pressure is

$$p = p_r + p_m + p_\varphi, \quad (29)$$

and satisfies the equation of state $p = \gamma(t) \rho$.

We find a time when the total pressure vanishes ($\xi = 1$)

$$a = \sqrt{\rho_{r0}/\rho_{\varphi0}} a_0. \quad (30)$$

When $\rho_r \ll \rho_m \ll \rho_\varphi$, which is valid enough at present, we get

$$\epsilon \equiv 2 - \xi \cong \frac{\rho_{m0}}{\rho_{\varphi0}} \frac{a_0}{a}. \quad (31)$$

Taking $|\omega| \sim 10^3$ [14] and $\eta \approx 3$ into account, we obtain $\epsilon \sim 10^{-3}$ from Eq. (9), and thus find $\rho_{m0}/\rho_{\varphi0} \sim 10^{-3}$ at present. If we adopt $\rho_{r0}/\rho_{m0} \sim 10^{-3}$, we get $a \sim 10^{-3} a_0$ for the time $\xi = 1$. It is interesting that it is the time when the hydrogen-recombination finished in the universe. As the critical density $\rho_c \sim 10^{-29} \text{ g.cm}^{-3}$ leads to the valid gravitational constant [10], we may adopt $\rho_{\varphi0} \sim \rho_c$. In this case, we find $\rho_{m0}/\rho_{\varphi0} \sim 10^{-2}$ and get $\epsilon \sim 10^{-2}$, which gives $|\omega| \sim 10^3$ with $\eta \approx 10$. When the expansion parameter $a(t) \rightarrow +\infty (t \rightarrow +\infty)$, the parameter ξ surely approaches to 2, the final state of the universe.

Our universe started (classically) from the Big Bang with $\xi = 0$ (the radiation era), passed the dust-dominated era ($\xi = 1$) rapidly in the early stage, and has been staying the negative pressure era ($\xi \approx 2$) for the almost all period of 10^{10} yr . The barotropic state of the universe has been varying extremely slowly from $\xi = 0$ to $\xi \approx 2$. At $t \sim 5 \times 10^9 \text{ yr}$, the parameter ϵ was 2×10^{-3} , and thus the gravitational constant was almost G_∞ as the present value is. This situation saves the evolution of life in the Earth in relation to the nuclear fusion of the Sun. Finally, the universe will approach to the state of $\xi = 2$ ($\gamma = -1/3$) as it expands for ever.

As a result of the barotropic evolution of matter in the universe,

the coupling function $\omega(\phi)$ diverges to the minus infinity when $t \rightarrow +\infty$ and the gravitational constant approaches dynamically to the constant G_∞ . The cosmological constant $\lambda(t)$ decreases in proportion to t^{-2} and converges to zero when $t \rightarrow +\infty$. The generalized scalar-tensor theory of gravitation reduces dynamically to general relativity. The Machian cosmological model approaches to the Friedmann universe in general relativity with $\ddot{a}(t) = 0$, $\lambda(t) = 0$, and $p = -\rho/3$.

The scalar field ϕ plays the most important role in the physical evolution of the universe and in the problem of dark matter. Our conjecture [11] on the barotropic history of the universe was proved. However, the next puzzle arises: Why does the scalar field ϕ have the exponential potential like Eq. (25)? We also conceive another conjecture: The coupling function ω is rather derived from the scalar field ϕ .

Acknowledgment

The author is grateful to Professor Carl Brans for helpful discussions and his hospitality at Loyola University (New Orleans) where this work was done. He would also like to thank the Nagasaki Prefectural Government for financial support.

References

- [1] C. Brans and R. H. Dicke, Phys. Rev. 124, 925 (1961).
- [2] N. Banerjee and S. Sen, Phys. Rev. D56, 1334 (1997).
- [3] V. Faraoni, Phys. Rev. D59, 084021 (1999).
- [4] A. Miyazaki, gr-qc/ 0012104, 2000
- [5] A. Miyazaki, gr-qc/ 0101112, 2001.
- [6] A. Miyazaki, gr-qc/ 0102003, 2001.
- [7] P. G. Bergmann, Int. J. Theor. Phys. 1, 25 (1968).

- [8] R. V. Wagoner, *Phys. Rev. D*1, 3209 (1970).
- [9] K. Nordtvedt, *Astrophys. J.* 161, 1059 (1970).
- [10] A. Miyazaki, *gr-qc/ 0102105*, 2001.
- [11] A. Miyazaki, *gr-qc/ 0103003*, 2001.
- [12] S. Perlmutter, M. S. Turner, and M. White, *Phys. Rev. Lett.* 83, 670 (1999).
- [13] T. Matos, F. S. Guzmán, and L. A. Ureña-López, *astro-ph/ 9908152*, 1999.
- [14] X. Chen and M. Kamionkowski, *Phys. Rev. D*60, 104036 (1999).
- [15] O. Bertolami and P. J. Martins, *Phys. Rev. D*61, 064007 (2000).